

# Universal features of black holes in the large $D$ limit

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# Why black hole dynamics is hard

*Non-decoupling:*

BH is an extended object whose dynamics mixes strongly with background

BH's own dynamics not well-localized, not decoupled

# Why black hole dynamics is hard

BHs, like other extended objects, have (quasi-) normal modes

but typically localized at some distance from the horizon

~ photon orbit in AF

in AdS backgrounds may be further away

→ hard to disentangle bh dynamics from background dynamics

# Why black hole dynamics is hard

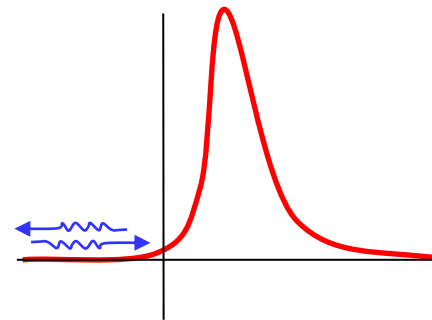
BH dynamics lacks a *generically small* parameter

Decoupling *requires* a small parameter

Near-extremality does it: AdS/CFT-type decoupling

Develop a throat

effective radial potential



# Large D limit

*Kol et al*

*RE+Suzuki+Tanabe*

1/D as small parameter

Separates bh's own dynamics from background spacetime

- *sharp* localization of bh dynamics

BH near-horizon well defined

- a very special  $2D$  bh

Somewhat similar to decoupling limit in ads/cft

# Large D limit

**Far-region:** background spacetime w/ holes

only knows bh size and shape

→ far-zone trivial dynamics

**Near-region:**

– non-trivial geometry

– large universality classes eg neutral bhs (rotating, AdS etc)

Large  $D$  expansion may help for

- **calculations**: new perturbative expansion
- deeper **understanding** of the theory (reformulation?)

**Universality** (due to strong localization)  
is good for both

# Large $D$ black holes

Basic solution

$$ds^2 = - \left( 1 - \left( \frac{r_0}{r} \right)^{D-3} \right) dt^2 + \frac{dr^2}{1 - \left( \frac{r_0}{r} \right)^{D-3}} + r^2 d\Omega_{D-2}$$

length scale  $r_0$



# Large $D$ black holes

$r_0$  **not** the only scale

Small *parameter*  $1/D \implies$  scale hierarchy

$$r_0/D \ll r_0$$

This is the **main feature** of large- $D$  GR

# Localization of interactions

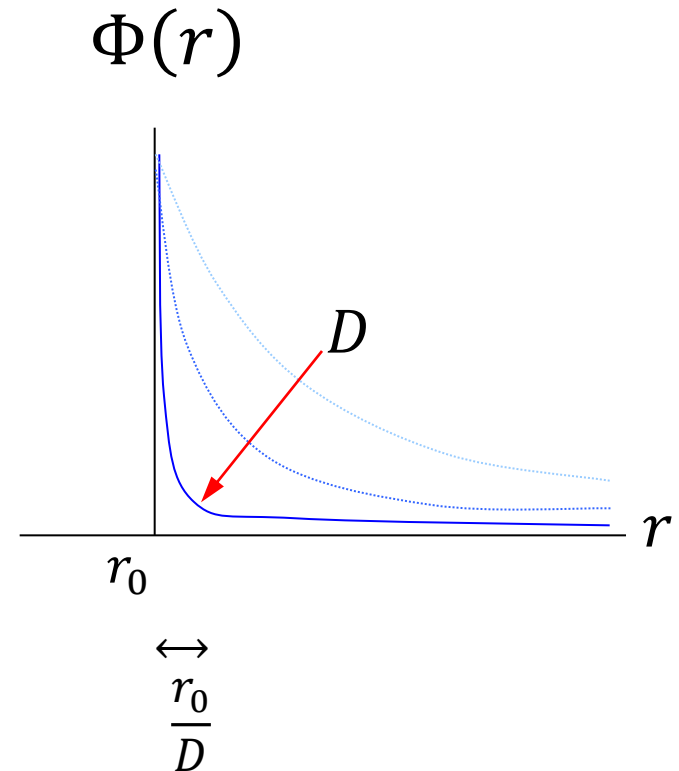
Large potential gradient:

$$\Phi(r) \sim \left(\frac{r_0}{r}\right)^{D-3}$$

$$\nabla\Phi \Big|_{r_0} \sim D/r_0$$

⇒ Hierarchy of scales

$$\frac{r_0}{D} \ll r_0$$



# Far zone

Fixed  $r > r_0$      $D \rightarrow \infty$

$$f(r) = 1 - \left(\frac{r_0}{r}\right)^{D-3} \rightarrow 1$$

$$ds^2 \rightarrow -dt^2 + dr^2 + r^2 d\Omega_{D-2}$$

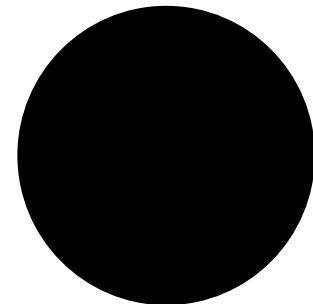
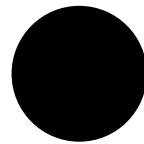
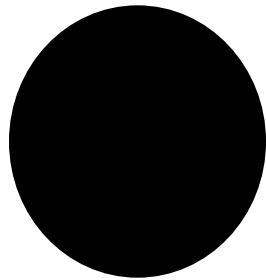
Flat, empty space at  $r > r_0$

no gravitational field

# *Far zone geometry*

scale  $\mathcal{O}(r_0 D^0)$

Holes cut out in Minkowski space



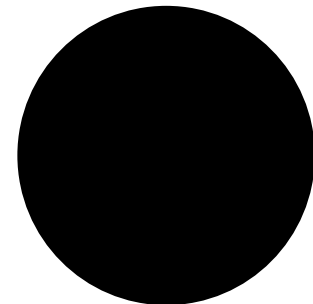
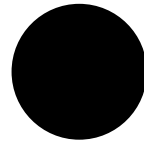
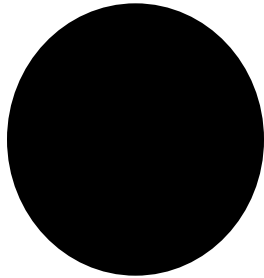
# Far zone

scale  $\mathcal{O}(r_0 D^0)$

Holes cut out in Minkowski space

No wave absorption (perfect reflection)

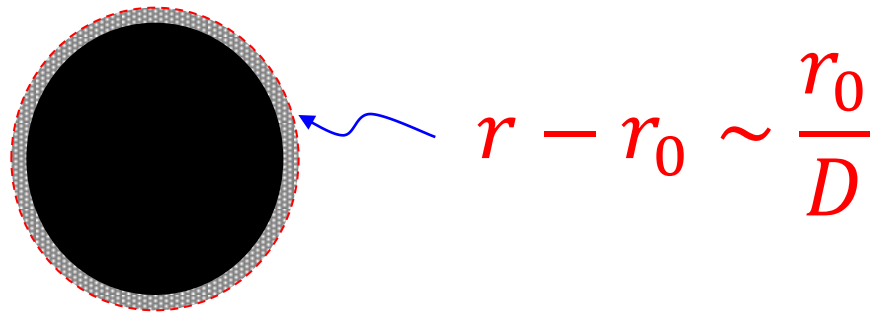
for  $D \rightarrow \infty$



# Near zone

Gravitational field appreciable only in *thin* near-horizon region

$$\left(\frac{r_0}{r}\right)^{D-3} = \mathcal{O}(1) \iff r - r_0 < \frac{r_0}{D}$$



# Near zone

Keep non-trivial gravitational field:

Length scales  $\sim r_0/D$  away from horizon

Surface gravity  $\kappa \sim D/r_0$  finite

Near-horizon coordinate:  $R = (r/r_0)^{D-3}$

All remain  $\mathcal{O}(1)$  where grav field is non-trivial

# Near zone

$$ds^2 = - \left( 1 - \left( \frac{r_0}{r} \right)^{D-3} \right) dt^2 + \frac{dr^2}{1 - \left( \frac{r_0}{r} \right)^{D-3}} + r^2 d\Omega_{D-2}$$

$$\left. \begin{aligned} \left( \frac{r}{r_0} \right)^{D-3} &= \cosh^2 \rho \\ t_{near} &= \frac{D}{2r_0} t \end{aligned} \right\} \begin{array}{l} \text{finite} \\ \text{as } D \rightarrow \infty \end{array}$$



# Near zone

$$dS_{nh}^2 \rightarrow \frac{4r_0^2}{D^2} \underbrace{(-\tanh^2 \rho dt_{near}^2 + d\rho^2)} + r_0^2 d\Omega_{D-2}^2$$

**2d string black hole**

*Elitzur et al  
Mandal et al  
Witten*

$$\ell_{string} \sim \frac{r_0}{D}, \quad \alpha' \sim \left(\frac{r_0}{D}\right)^2$$

*Soda  
Grumiller et al*

# *Near zone universality: neutral bhs*

2d string bh is near-horizon geometry  
of **all neutral non-extremal bhs**

- **rotation** appears as a **local boost**  
(in a third direction)
- **cosmo const shifts** 2d bh **mass**

More near-horizon structure than just  
Rindler limit

# *Near zone universality*

Charge modifies near-horizon geom  
some are 'stringy' bhs

eg, 3d black string *Horne+Horowitz*

but many different solutions possess  
same near-horizon

*universality classes*

## Large D expansion:

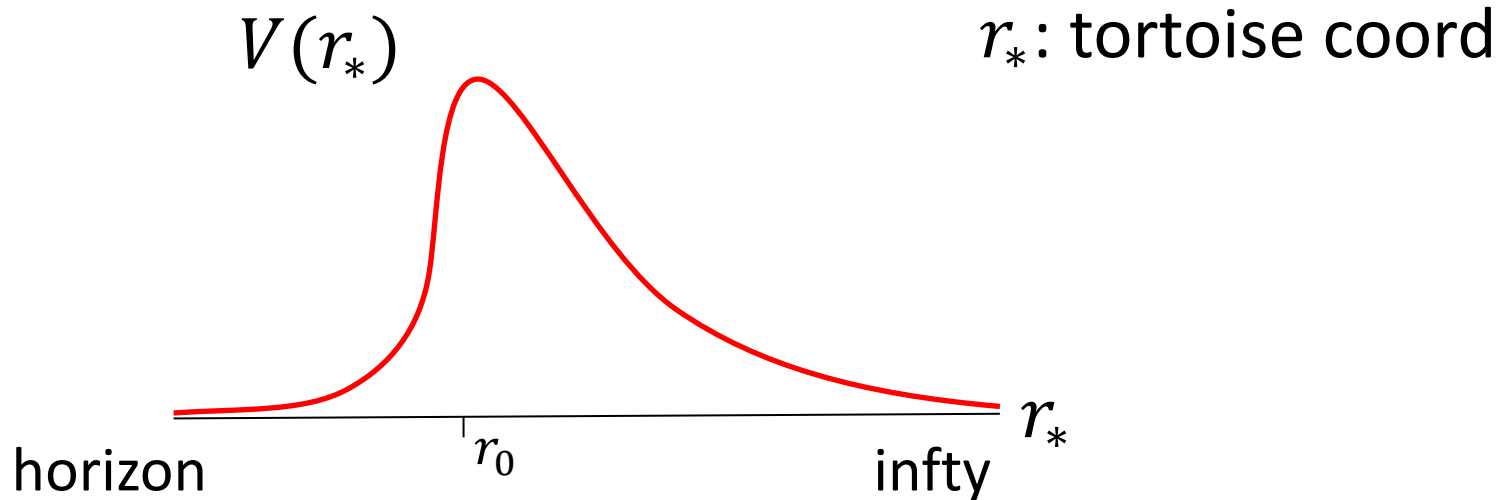
1. BH quasinormal modes
2. Instability of rotating bhs

# Massless scalar field

$$\square\Phi = 0$$

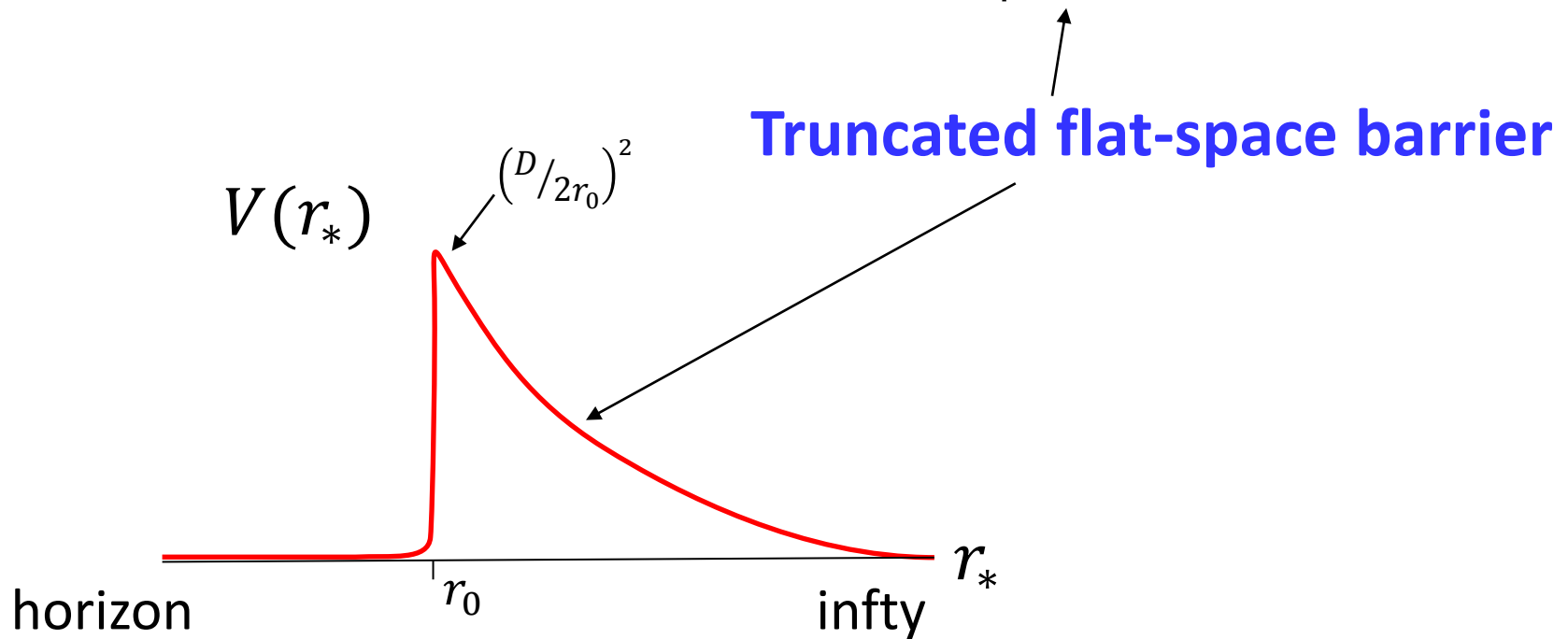
$$\Phi = r^{-\frac{D-2}{2}} \phi(r) e^{-i\omega t} Y_\ell(\Omega)$$

$$\frac{d^2\phi}{dr_*^2} + (\omega^2 - V(r_*))\phi = 0$$



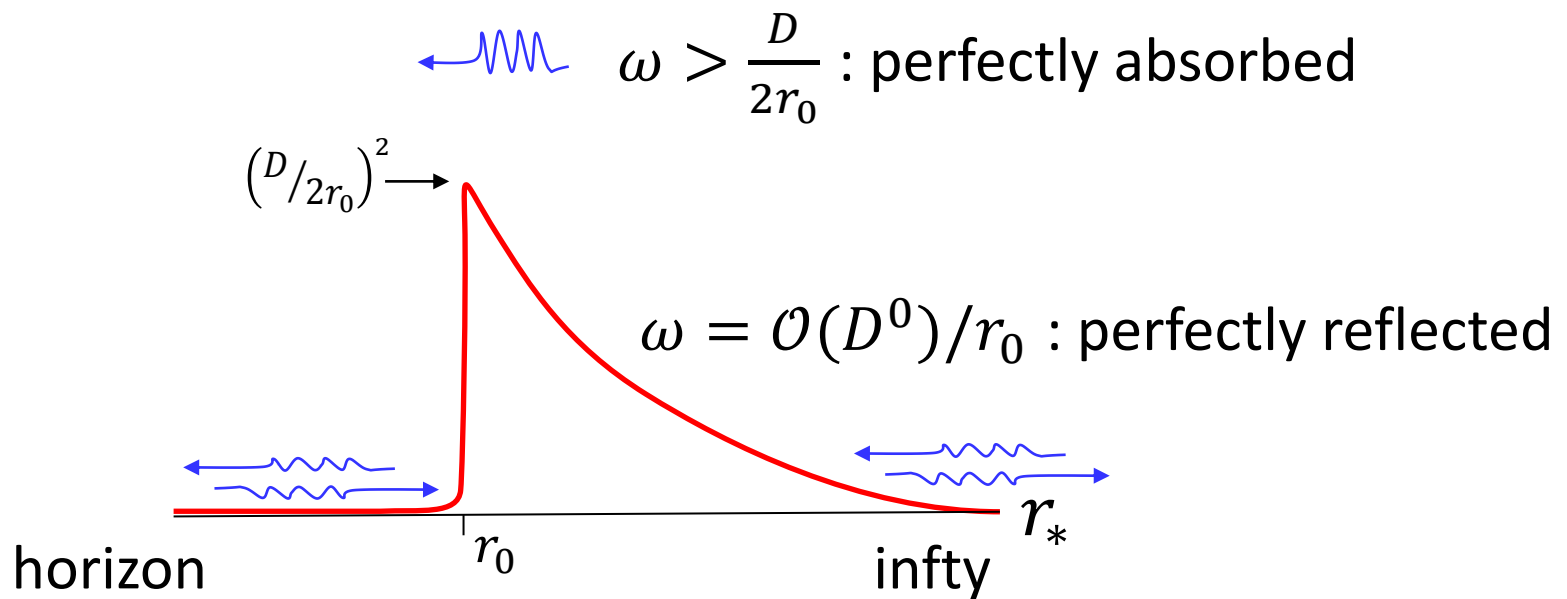
# Massless scalar field

$$D \rightarrow \infty \quad V(r_*) \rightarrow \frac{D^2}{4r_*^2} \Theta(r_* - r_0)$$



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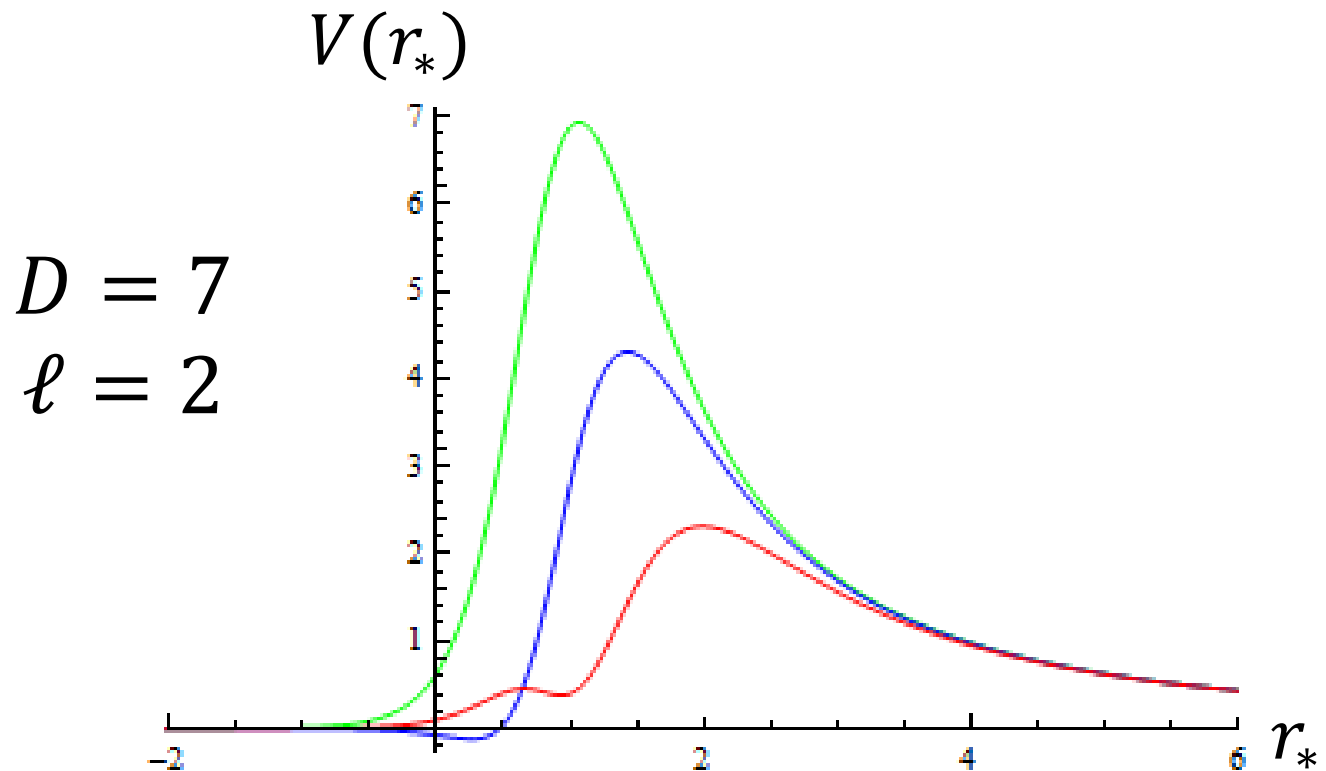


# Schwarzschild bh grav perturbations

*Kodama+Ishibashi*

Gravitational **scalar**, **vector**, **tensor** modes

$SO(D - 1)$  reps



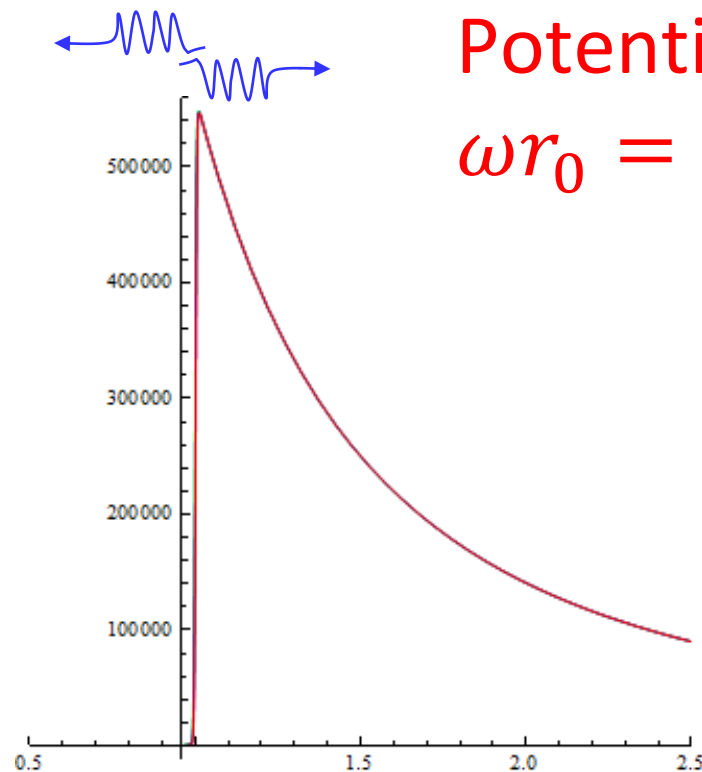


# Schwarzschild bh grav perturbations

scalar vector tensor

$$D = 500$$

$$\ell = 500$$



# Schwarzschild bh grav perturbations

scalar vector tensor

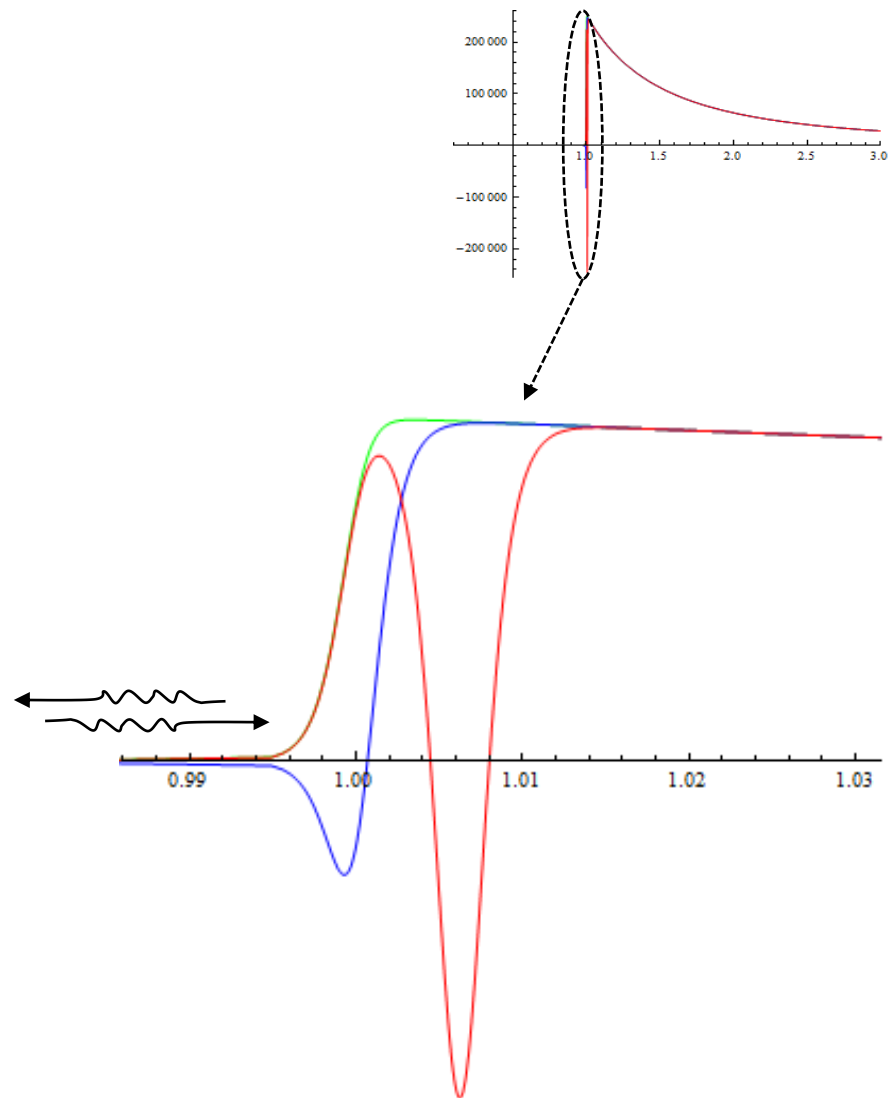
Potential seen by

$$\omega r_0 = \mathcal{O}(1)$$

$$\ell = \mathcal{O}(1)$$

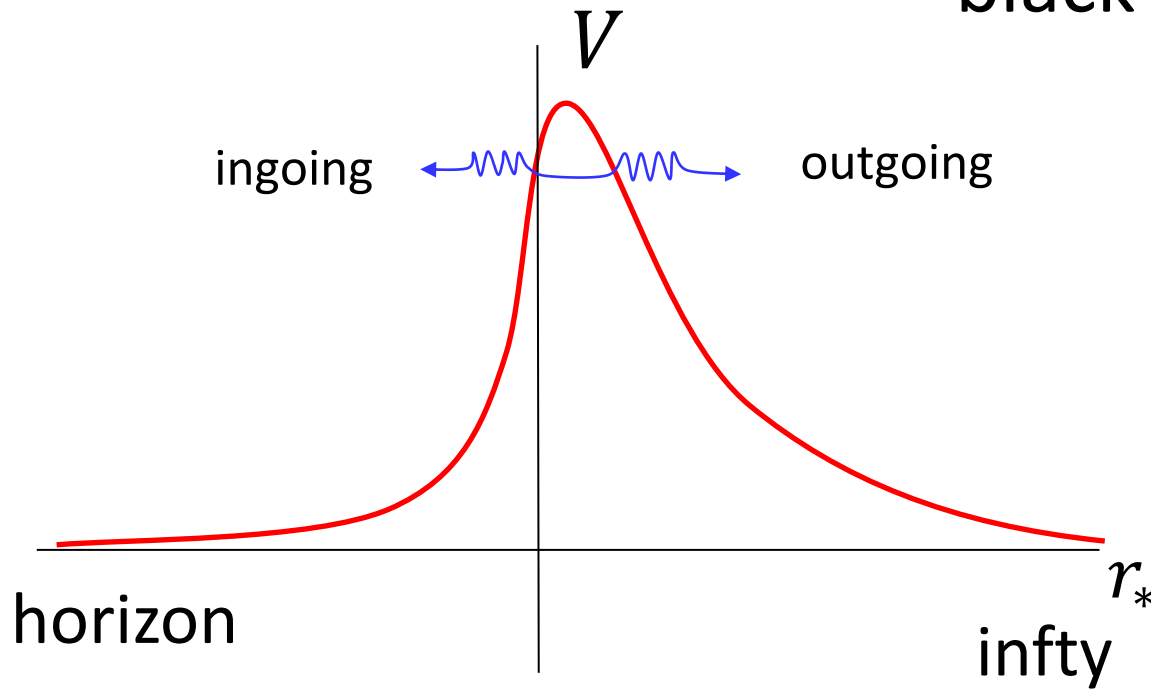
$$D = 1000$$

$$\ell = 2$$



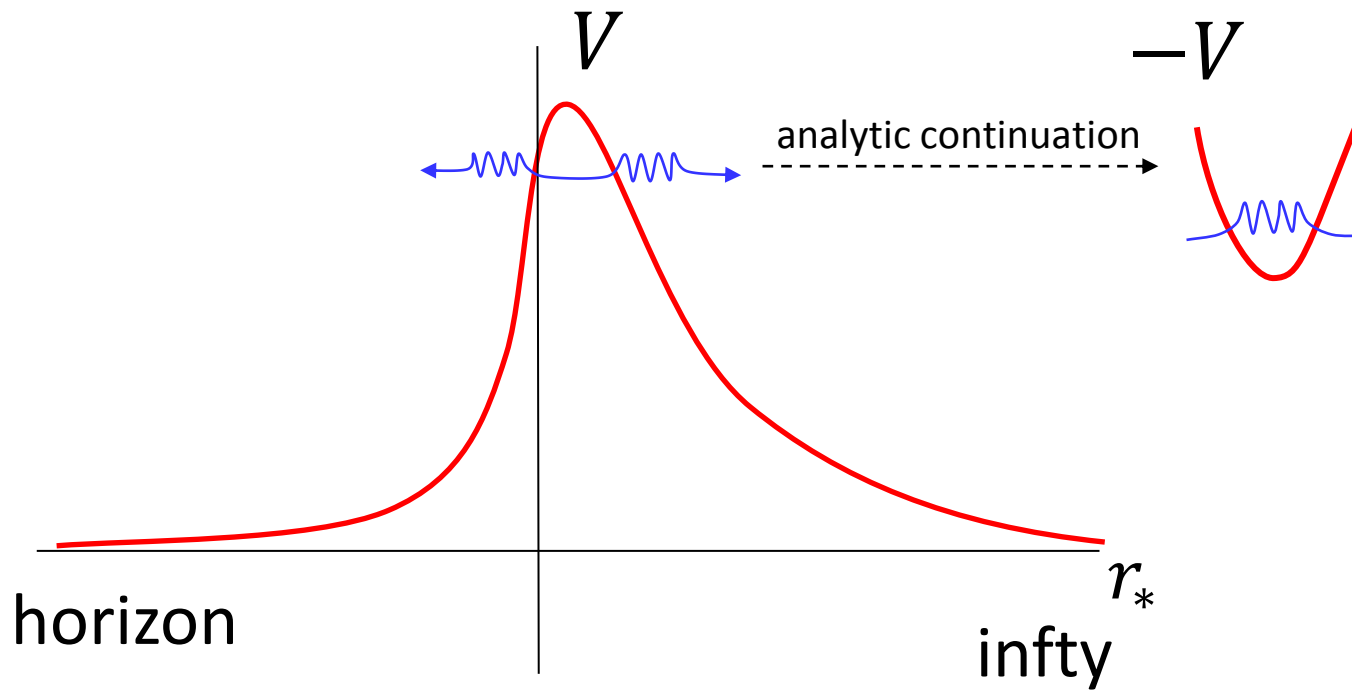
# Quasinormal modes

Free, damped oscillations of black hole



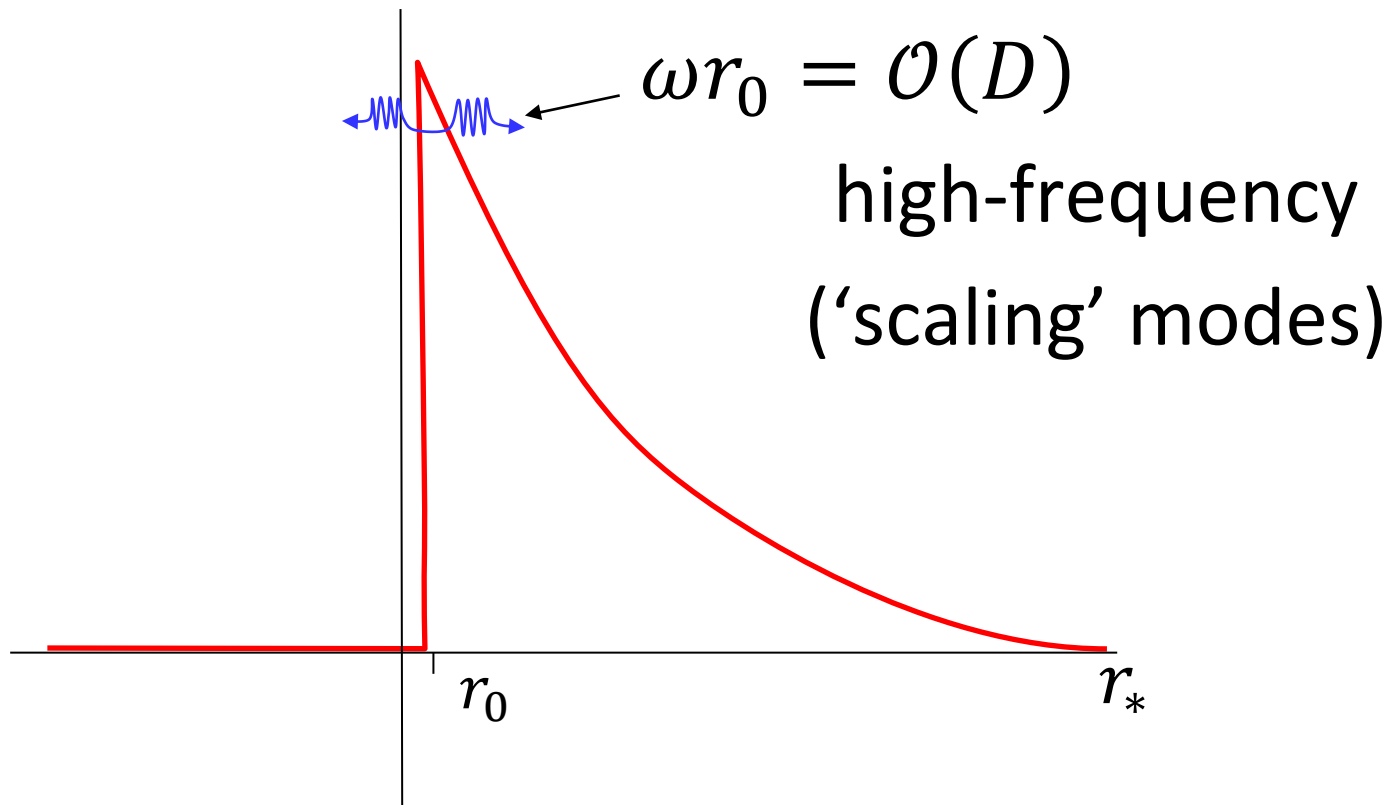
# Quasinormal modes

QNMs as bound states in inverted potential



$\omega r_0 = \mathcal{O}(D)$  QNMs

$$V(r_*) \rightarrow \frac{D^2}{4r_*^2} \Theta(r_* - r_0)$$



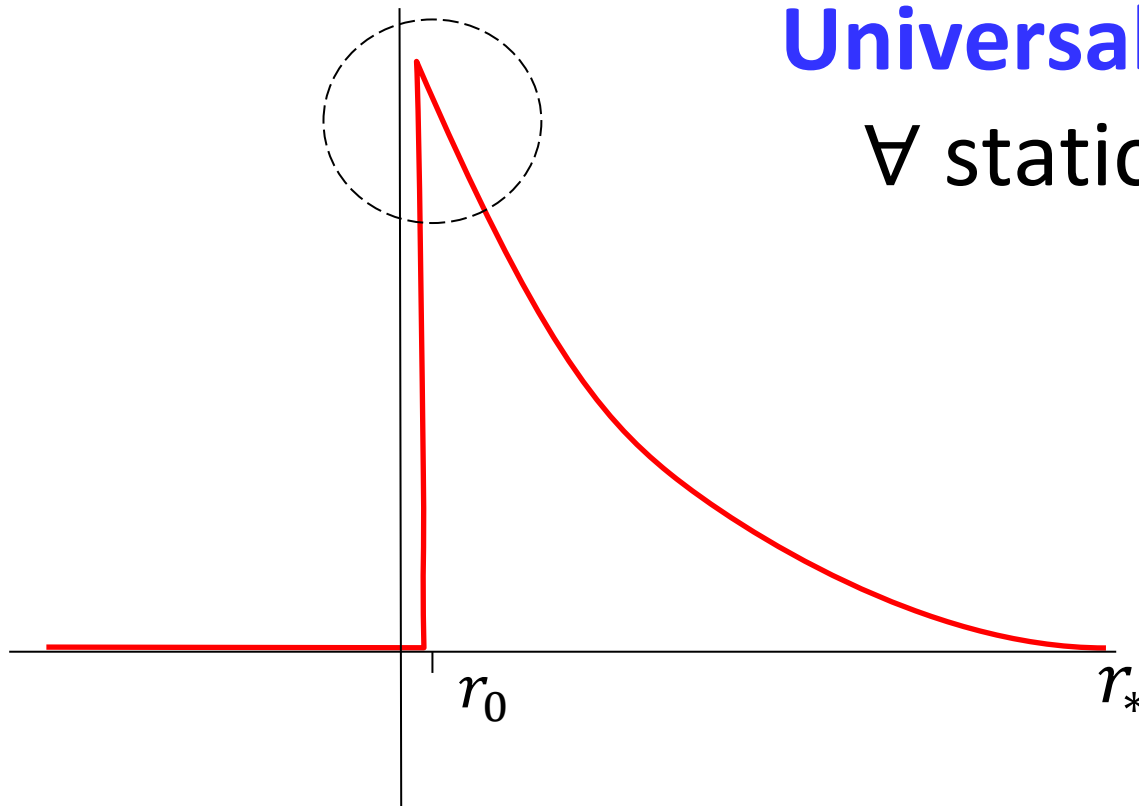
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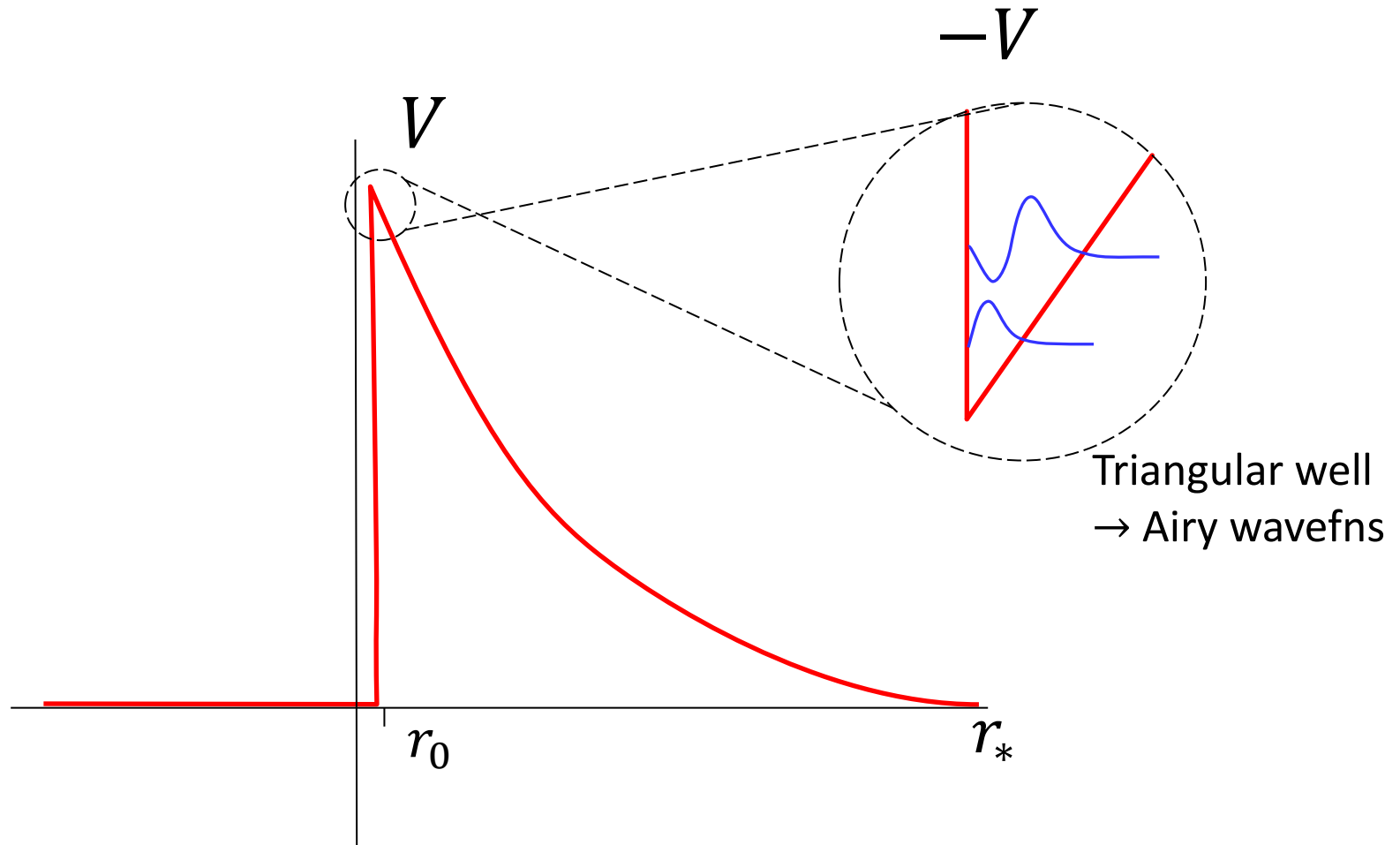
**Holes in flat space**

**Universal structure**

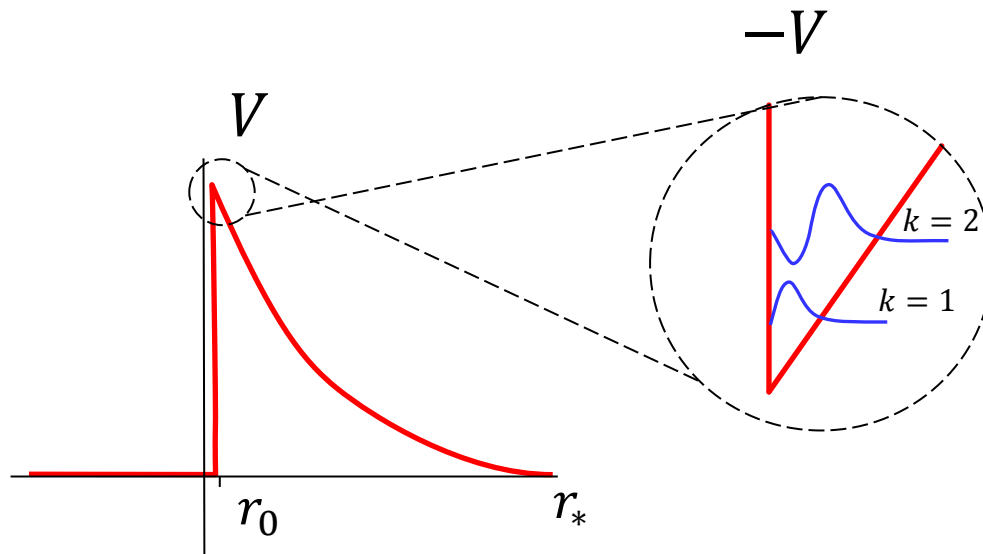
$\forall$  static, AF bhs



$$\omega r_0 = \mathcal{O}(D) \text{ QNMs}$$



$\omega r_0 = \mathcal{O}(D)$  QNMs



$$\Rightarrow \omega_{(\ell, k)} r_0 = \frac{D}{2} + \ell - \left( \frac{e^{i\pi}}{2} \left( \frac{D}{2} + \ell \right) \right)^{\frac{1}{3}} a_k$$

Airy zeroes  $\swarrow$



# Universal spectrum @ large D

$$\omega_{(\ell,k)} r_0 = \frac{D}{2} + \ell - \left( \frac{e^{i\pi}}{2} \left( \frac{D}{2} + \ell \right) \right)^{\frac{1}{3}} a_k$$

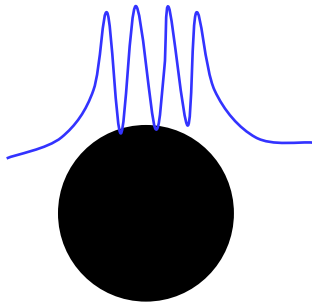
Depends **only on** bh radius  $r_0$

**Same** spectrum for:

- any charges, dilaton coupling etc
- scalar, vector, tensor perturbations

# Universal spectrum @ large D

$$\omega_{(\ell,k)} r_0 = \frac{D}{2} + \ell - \left( \frac{e^{i\pi}}{2} \left( \frac{D}{2} + \ell \right) \right)^{\frac{1}{3}} a_k$$



spectrum of scalar  
**oscillations of a hole**  
in space

$\frac{\text{Im}\omega}{\text{Re}\omega} \sim D^{-2/3} \rightarrow 0$ : sharp resonances  
'normal modes' of bh

$\omega r_0 = \mathcal{O}(1)$  QNMs

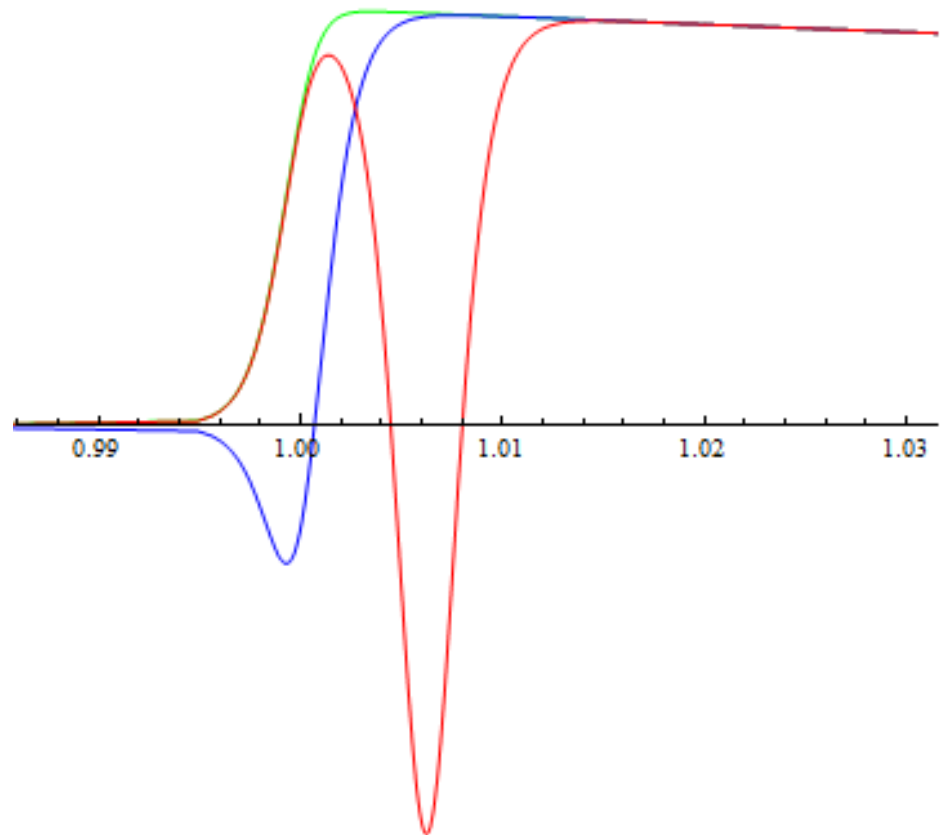
More complicated  
wave eqn

but we've solved it  
up to

$D^{-3}$  for vectors

$D^{-2}$  for scalars

(no tensors)



# Quantitative accuracy

$\omega r_0 = \mathcal{O}(1)$  modes

Vector mode (purely imaginary)

- At  $D = 100$ :

$\ell = 2$  mode     $\text{Im } \omega r_0 = -1.01044742$  (analytical)

-1.01044741 (numerical *Dias et al*)

# Quantitative accuracy

$\omega r_0 = \mathcal{O}(1)$  modes

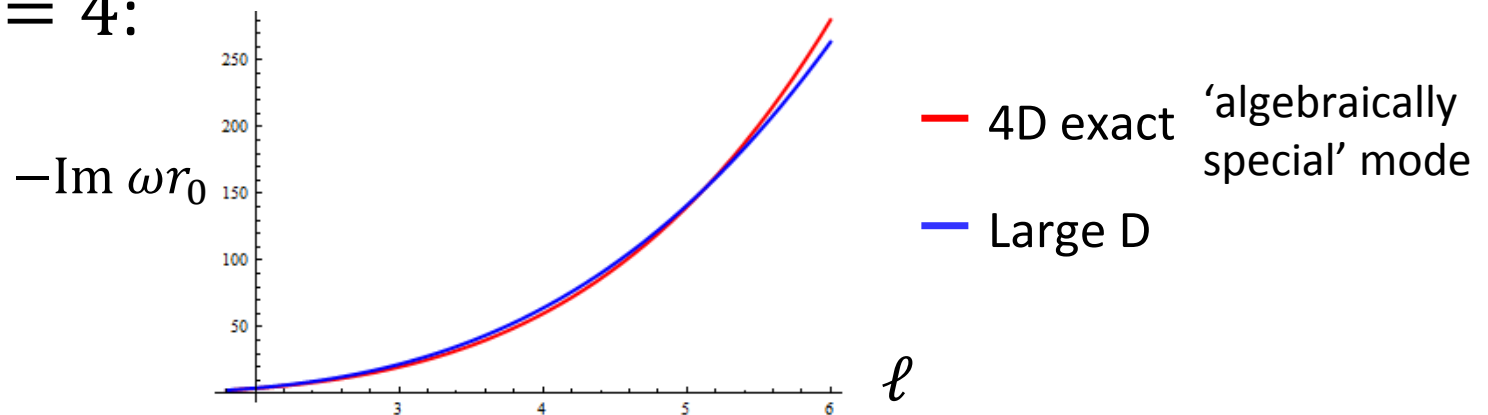
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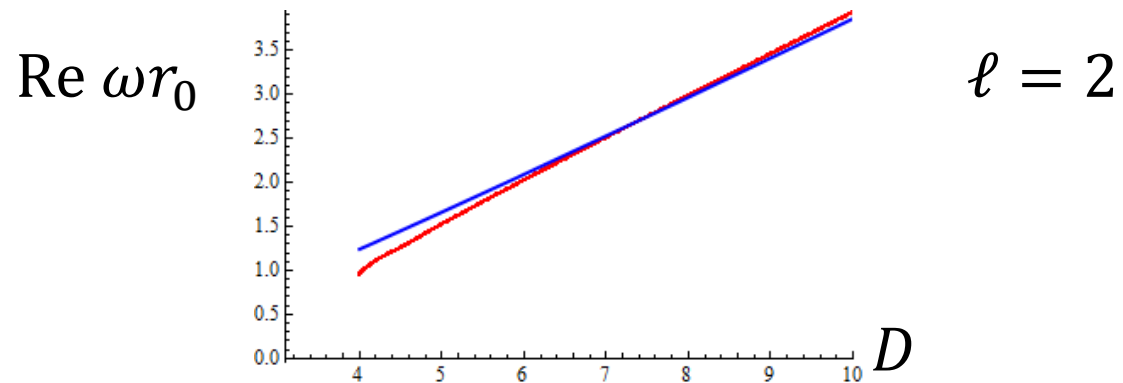
- At  $D = 4$ :



# Quantitative accuracy

$\omega r_0 = \mathcal{O}(D)$  modes

$\text{Re } \omega r_0 = \frac{D}{2} + \ell$  : good at moderate  $D$



$\text{Im } \omega r_0 \sim D^{1/3}$  : only good at *very* high  $D$

$D \gtrsim 300$  (!)

Instability of rotating bhs

Hi-D bhs have *ultra-spinning* regimes

Expect instabilities:

- axisymmetric
- non-axisymmetric (at lower rotation)

Confirmed by numerical studies *Dias et al*

*Hartnett+Santos*

*Shibata+Yoshino*

Analytically solvable in  $1/D$  expansion thanks to universality features – *also in AdS*



# Equal-spin, odd-D, Myers-Perry black holes

→ only radial dependence → ODEs

But equations are coupled

– analytically hopeless



Equations *do decouple* for rotation=0

## Large D expansion:

Leading large D near-horizon: rotating bh is just a **boost of Schw**

→ rotating eqns decouple

can be solved analytically

Beyond leading order, MP metric is not boosted Schw, but LO boost allows to decouple eqns

## Analytical computation of QNMs

- *Axisymmetric* instability for

$$a > \frac{\sqrt{3}}{2} r_+$$

- *Non-axisymmetric* instability for

$$a > \frac{1}{\sqrt{2}} r_+ = .71 r_+$$

*Comparison to numerical:*

$$D=5: a > .81r_+ \quad , \quad D=15: a > .73r_+ \quad \text{Hartnett+Santos}$$

# Outlook

*Any* problem that can be formulated in  
arbitrary  $D$  is amenable to **large  $D$**   
**expansion**

simpler, even analytically solvable

## Universal features

**Far:** *empty space*  $\forall bhs$

**Near:** *2D string bh*  $\forall neutral bhs$

BH dynamics splits into:

$\omega r_0 = \mathcal{O}(D)$  : **non-decoupled** modes

- scalar field **oscillations of a hole** in space
- universal *normal* modes

$\omega r_0 = \mathcal{O}(D^0)$  : **decoupled** modes

- localized in near-horizon region